Advanced LIGO:

Next Steps and Future Improvements

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LIGO-G1300351

LVC meeting, Washington DC, March 2013

Contents

- Background / Introduction
- A Zoo of technologies to reduce individual noise sources.
- Example upgrade 1: Team Red design
- Example upgrade 2: Team Green design(s)
- Example upgrade 3: Team Blue design
- Example upgrade 4: Team Red Xylophone
- What can we learn from all this?

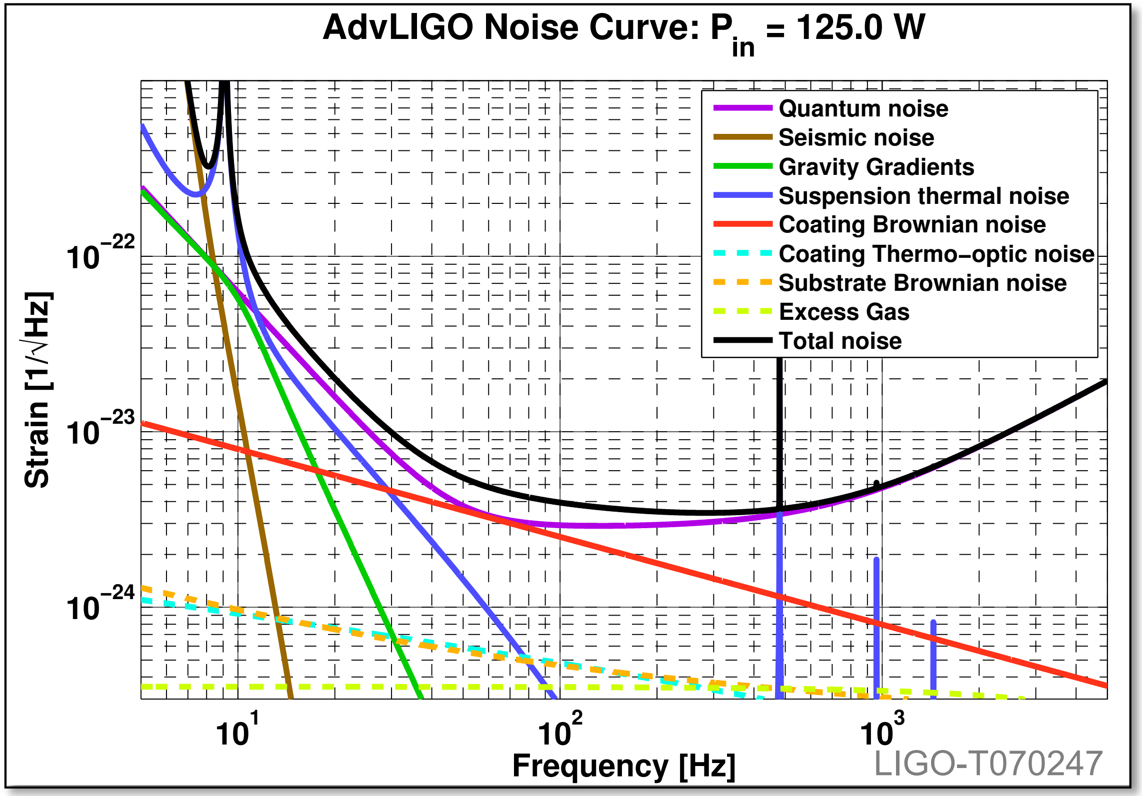
Increasing Effort, Cost, Ambition

Motivation for Advanced LIGO upgrades

- The advanced LIGO baseline detectors are expected to accomplish the first direct detection of gravitational waves.
  - See also Abadie et al, CQG, **2010**, *27*, 173001
- However, these observations are likely to be of modest signal-to-noise ratio (SNR). If we want to access the full physics of the sources we will need to increase the SNR.
- As we will see it seems possible to upgrade the aLIGO instruments gaining a broadband sensitivity improvement by a factor of 3-5 (roughly equivalent to increasing the event rate by a factor 25-100).
- For details on the exciting science aLIGO upgrades will bring into our reach please see: Adhikari et al: 'Astrophysical Motivations for the Third Generation LIGO Detectors', **LIGO-T1200099–v2**

Noise Sources limiting the Advanced Detectors

- In order to understand how we can potentially improve 2G detectors, we need to see what they are limited by:

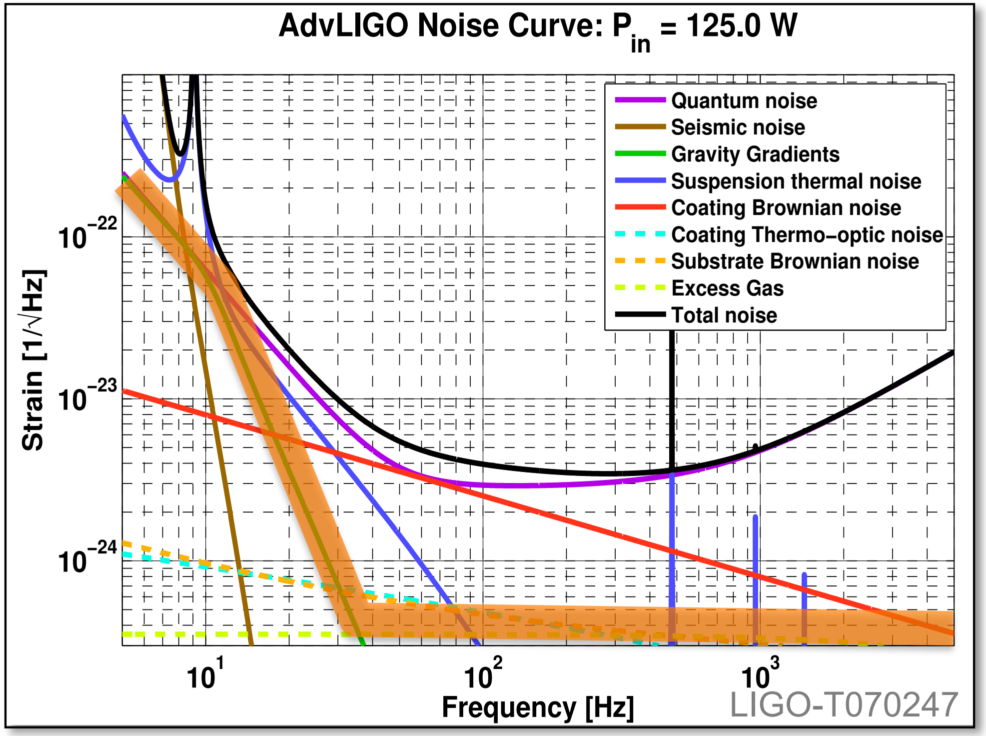


- **Quantum Noise** limits most of the frequency range.
- **Coating Brownian** limits (or is close) in the range from 50 to 100Hz.
- Below 50Hz we are limited by 'walls' made of **Suspension Thermal**, **Gravity Gradient** and **Seismic noise**.

Slide 4

Upgrades within the Advanced LIGO infrastructure

- The advanced LIGO baseline sensitivity is far away from the infrastructure limits.
- Infrastructure limit is usually defined as combination of residual gas noise and gravity gradient noise.
- **So there is plenty of room for advanced LIGO upgrades within the existing infrastructure! And this will be the focus of the rest of this presentation.**



Strawman Exercise

- About 1.5 years ago the LIGO Scientific Collaboration decided to initiate an effort to develop simple design studies (so-called Strawman designs) for aLIGO upgrades.
- 3 teams formed: Blue (headed by R.Adhikari), Green (headed by S.Ballmer) and Red (headed by S.Hild).
- Some interesting aspects:

When do we need to be ready for the upgrades?

What R&D do we need to carry out over the next few years?

What is the available budget?

Can we do the upgrades in an incremental way?

How much improvent is possible?

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Mirror Thermal Noise

- Due to thermal fluctuations the position of the mirror sensed by the laser beam is not necessarily a good representation of the center of mass of the mirror.
- Various noise terms involved: Brownian, thermo-elastic and thermo-refractive noise of each substrate and coating (or coherent combinations of these, such as thermo-optic noise).
- For nearly all current and future designs coating Brownian is the dominating noise source:

$$S_x(f) = \frac{4k_\mathrm{B}T}{\pi^2 f Y}\frac{d}{r_0^2}\left(\frac{Y'}{Y}\phi_{\|} + \frac{Y}{Y'}\phi_{\perp}\right)$$

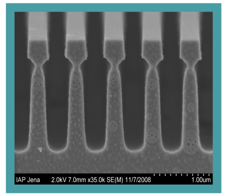
PSD of displacement ($S_x(f)$); Boltzmann constant (k_B); Temperature (T); Geometrical coating thickness (d); Loss angle of coating ($\phi_{\|}$, $\phi_{\perp}$); Young's modulus of mirror substrate (Y); laser beam radius (r_0); Young's modulus of coating (Y')

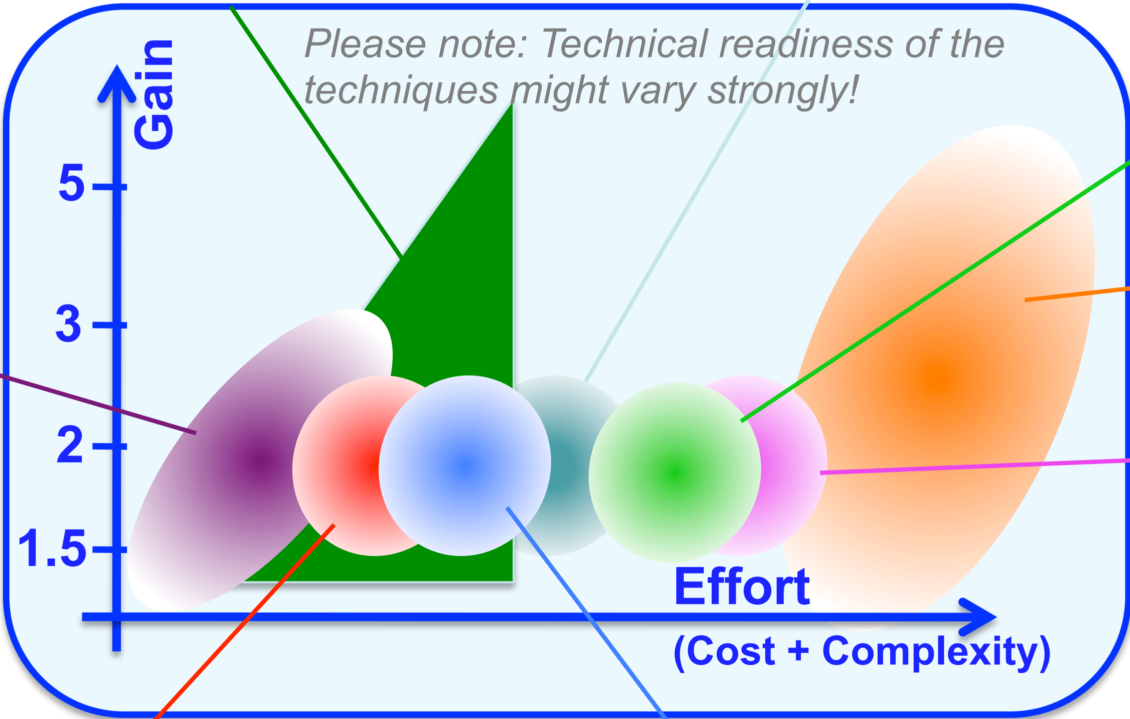
Harry et al, CQG 19, 897–917, 2002

How to reduce Mirror Thermal Noise?

Improved coating materials (e.g. crystalline coatings like AlGaAs, GaPAs)
Cole et al, APL 92, 261108, 2008

Waveguide mirrors
Brueckner et al, Opt. Expr 17, 163, 2009
Heinert et al LIGO P1300034-v1



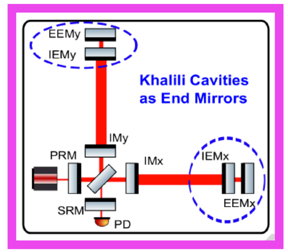


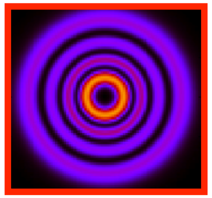
Resonant Delay Line

Cryogenic mirrors
Uchiyama et al, PRL **108**, 141101 (2012)

Larger beam size (needs larger mirrors)
Harry et al, CQG 19, 897–917, 2002

Khalili cavities
Khalili, PLA 334, 67, 2005
Gurkovsky et al, PLA 375, 4147, 2011





Different beam shape
Mours et al, CQG, 2006, 23, 5777
Chelkowski et al, PRD, 2009, 79, 122002

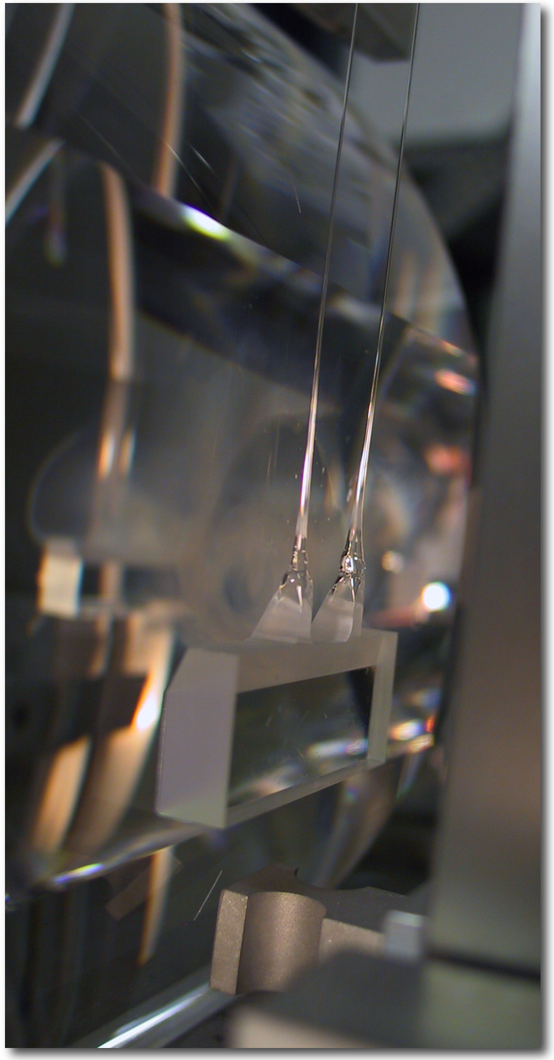
Amorphous Silicon coatings
Liu et al, PRB 58, 9067, 1998

Suspension Thermal Noise

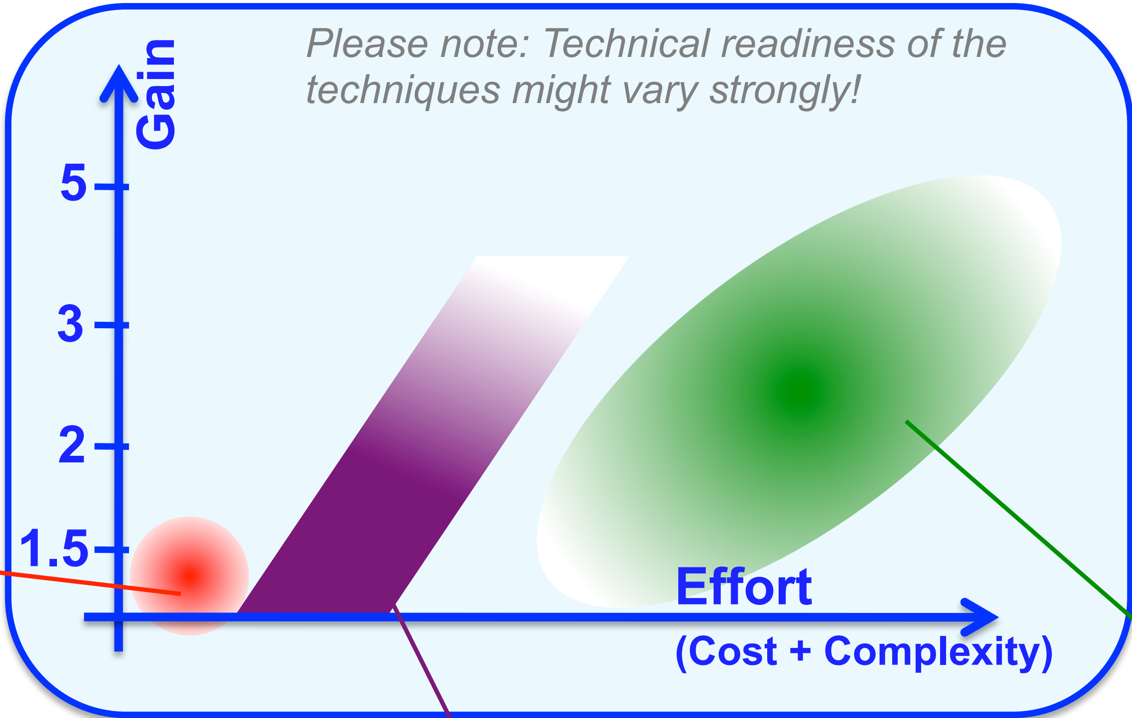
Boltzmann constant · Temperature · Loss angle · PSD of displacement · Mirror mass · Resonance frequency

$$x^2(\omega) = \frac{4k_{\mathrm{B}}T\omega_0^2\phi(\omega)}{\omega m[(\omega_0^2-\omega^2)^2+\omega_0^4\phi^2(\omega)]}$$

- Mirrors need to be suspended in order to decouple them from seismic.
- Thermal noise in metal wires and glass fibres causes horizontal movement of mirror.
- Relevant loss terms originate from the bulk, surface and thermo-elastic loss of the fibres + bond and weld loss.
- Thermal noise in blade springs causes vertical movement which couples via imperfections of the suspension into horizontal noise.



How to reduce Suspension Thermal Noise?



Improve fibre geometry/ profile

Bending points, energy stored via bending and neck profile can be potentially further optimised.

Increase length of final pendulum stage.

Allows the push suspension thermal noise out detection band.

Cooling of the suspension to cryogenic temperatures.

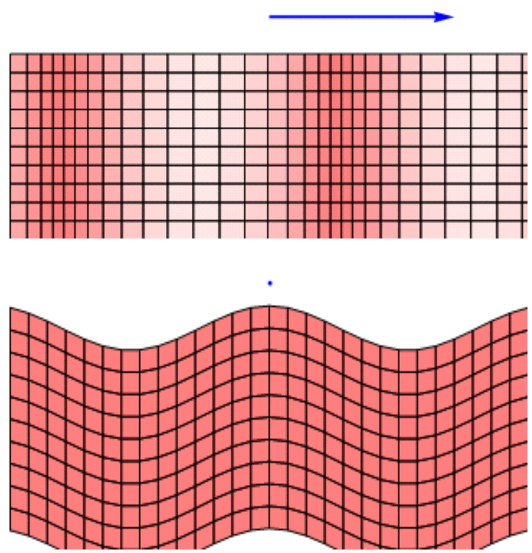
Usually also requires a change of materials.

Gravity Gradient Noise (also referred to as Newtonian noise)

- Seismic causes density changes in the ground and shaking of the mirror environment (walls, buildings, vacuum system).
- These fluctuations cause a change in the gravitational force acting on the mirror.
- Cannot shield the mirror from gravity. ☹

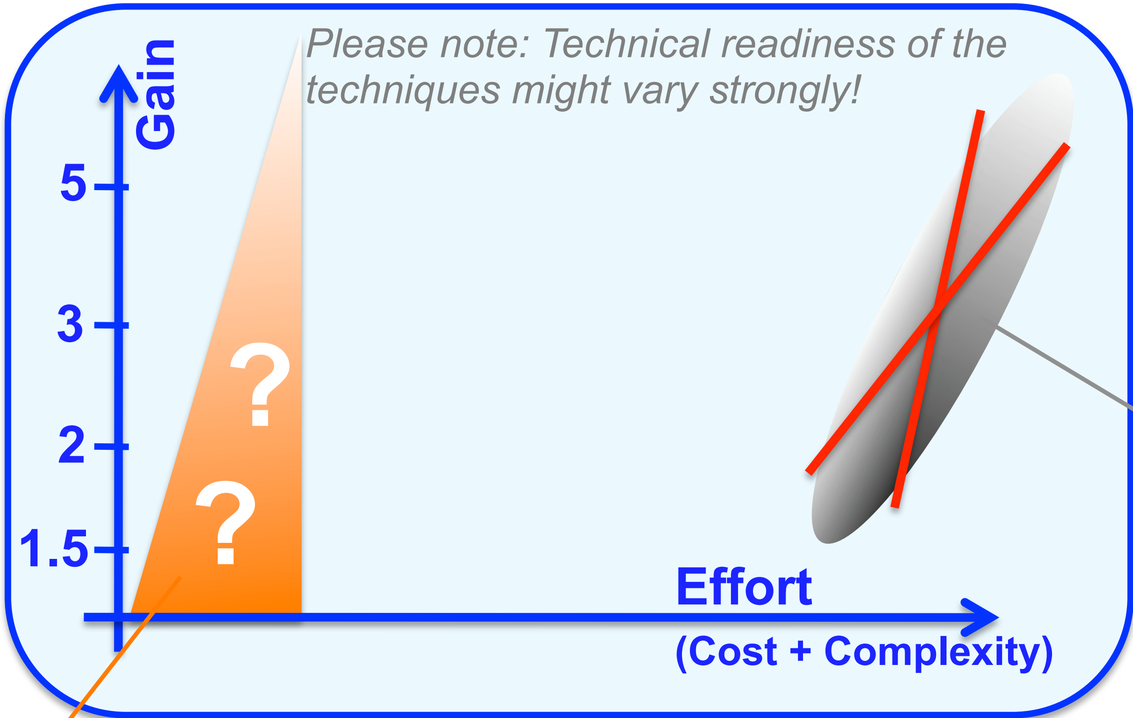
$$N_{\mathrm{GG}}(f)^2 = \frac{4\cdot\beta^2\cdot G^2\cdot\rho_r^2}{L^2\cdot f^4}\cdot X_{\mathrm{seis}}^2,$$

- PSD of strain: $N_{\mathrm{GG}}(f)^2$
- Coupling constant (depends on type of seismic waves, soil properties, etc): β
- Gravitational constant: G
- Density of ground: ρ_r
- Arm length: L
- frequency: f
- PSD of seismic: X_{seis}



Images: courtesy G.Cella

How to reduce Gravity Gradient Noise?



Reduce seismic noise at site., i.e. select a quieter site, potentially underground.

Beker et al, Journal of Physics: Conference Series 363 (2012) 012004

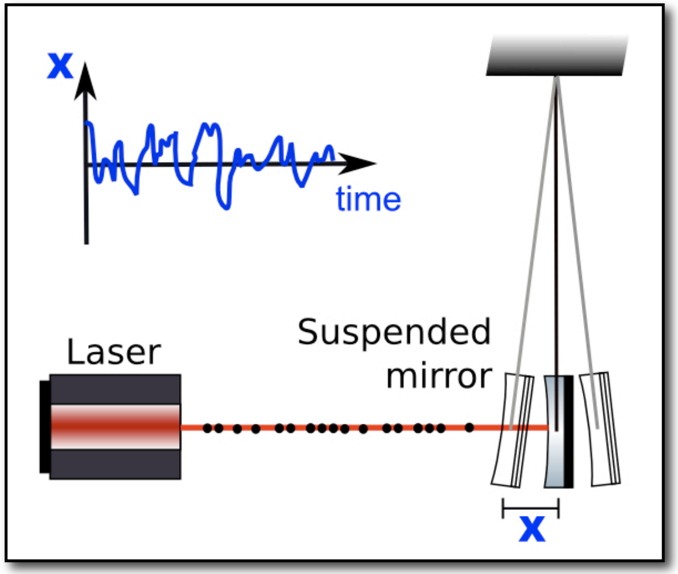
Obviously not possible within the LIGO infrastructure (but consider for other projects, see GW4 session tomorrow)

Subtraction of gravity gradient noise using an array of seismometers.

- Beker et al: General Relativity and Gravitation Volume 43, Number 2 (2011), 623-656
- Driggers et al: arXiv:1207.0275v1 [gr-qc]

Quantum Noise

- Quantum noise is a direct manifestation of the **Heisenberg Uncertainty Principle**.
- It is comprised of **photon shot noise (sensing noise)** at high frequencies and **photon radiation pressure noise (back-action noise)** at low frequencies.



$$h_{\mathrm{sn}}(f) = \frac{1}{L}\sqrt{\frac{\hbar c \lambda}{2\pi P}}$$

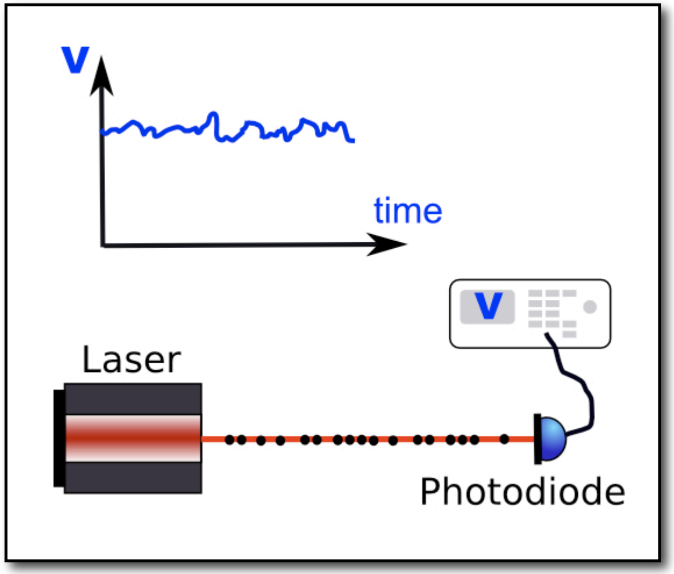
$$h_{\mathrm{rp}}(f) = \frac{1}{m f^2 L}\sqrt{\frac{\hbar P}{2\pi^3 c \lambda}}$$

wavelength

optical power

Arm length

Mirror mass



photon radiation pressure noise

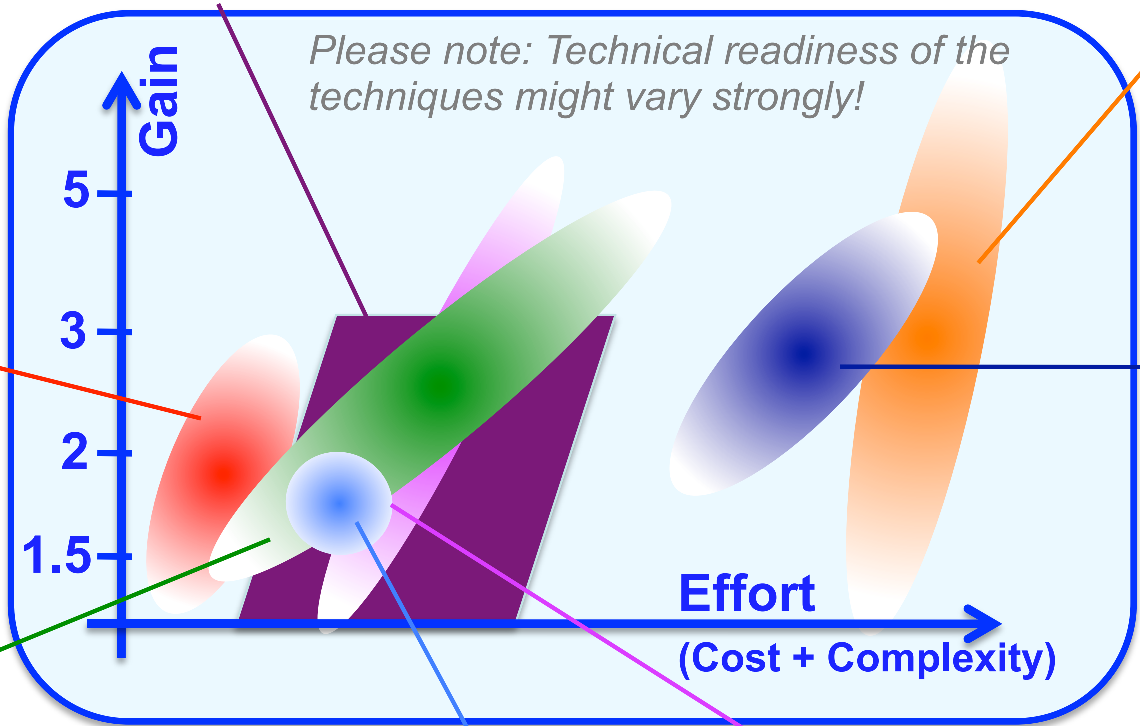
photon shot noise

How to reduce Quantum Noise?

Squeezing with frequency dependent squeezing angle
Kimble et al, PRD 65, 2002

Speedmeter
Measures momentum of test masses and is therefore not susceptible to Heisenberg Uncertainity Principle.
Chen, PRD 67, 122004, 2003

Squeezed Light
LIGO Scientific collaboration, Nature Phys. 7 962–65, 2011



Optical Bar + Optical Lever
Khalili, PLA 298, 308-14, 2002

Increased Laser Power
Need to deal with thermal problems and instabilities

Local readout
Rehbein et al, PRD 78, 062003, 2008

Increased Mirror Weight
Need to deal with thermal problems and instabilities

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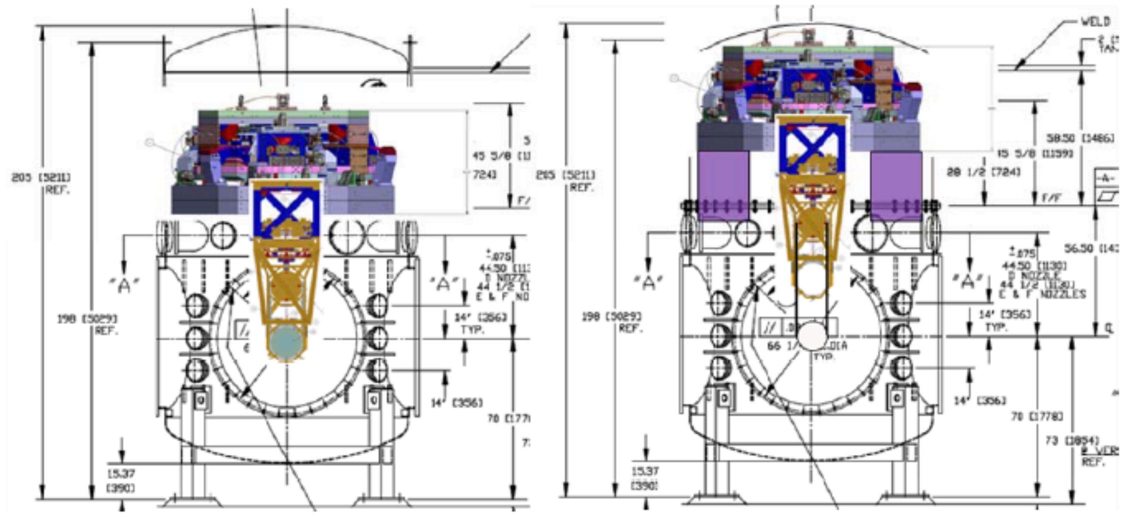
Increasing Effort, Cost, Ambition

Suspension Thermal Noise

Assume a boosted aLIGO Quad-suspension:

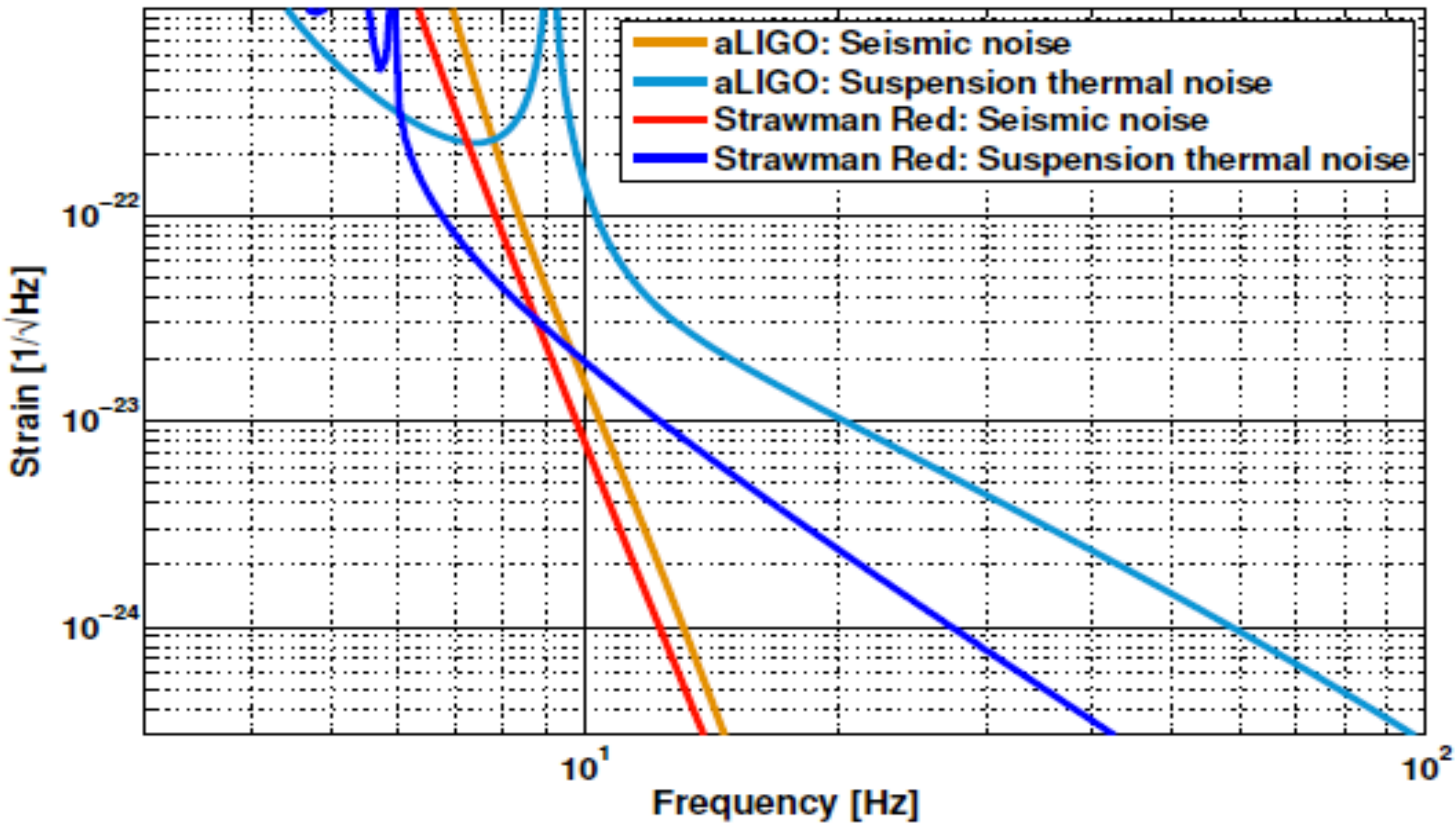
- Increased length of last stage to **1.2m** to reduce suspension thermal noise.
- Increased **mirror mass of 160kg** to reduce suspension thermal noise (and radiation pressure noise and coating noise)



(a) aLIGO QUAD (b) LIGO-3 QUAD

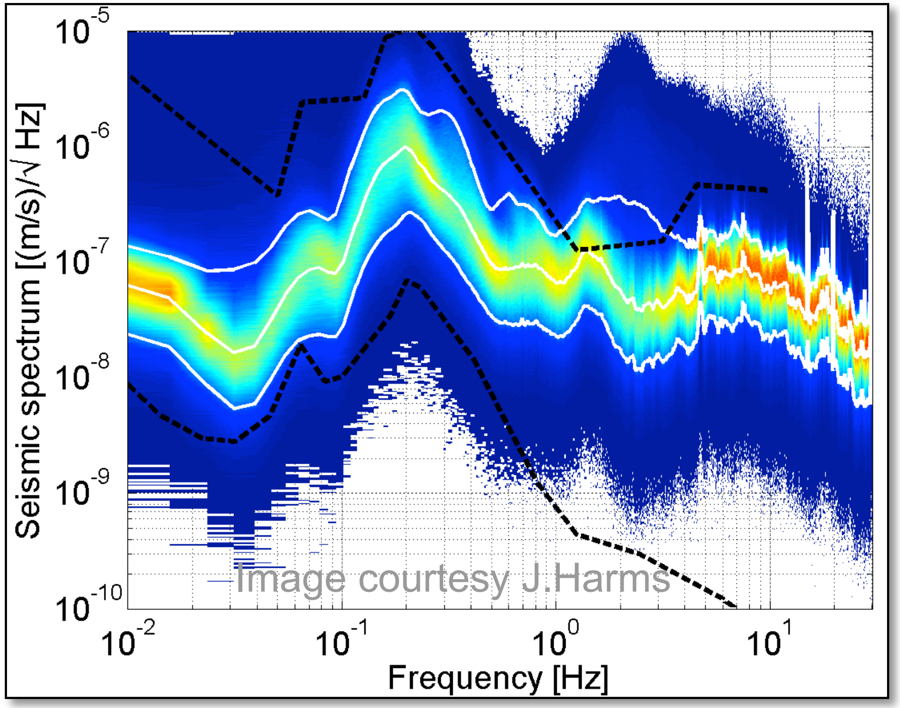
Test Masses and Suspensions		
Mirror Material	Fused Silica	Fused Silica
Main Test Mass Diameter	35 cm	55 cm
Main Test Mass Weight	42 kg	160 kg
Masses in Main Quad (from top)	22 kg/22 kg/40 kg/40 kg	44 kg/66 kg/120 kg/160 kg
Masses in Reaction Chain (from top)	22 kg/22 kg/40 kg/40 kg	22 kg/22 kg/40 kg/40 kg
Total Mass of a Main Suspension	250 kg	520 kg
Length of Final Suspension Stage	0.6 m	1.2 m
Fused Silica Fibre Diameter	400 μm	566 μm
Fibre Diameter at Bending Point	800 μm	1624 μm

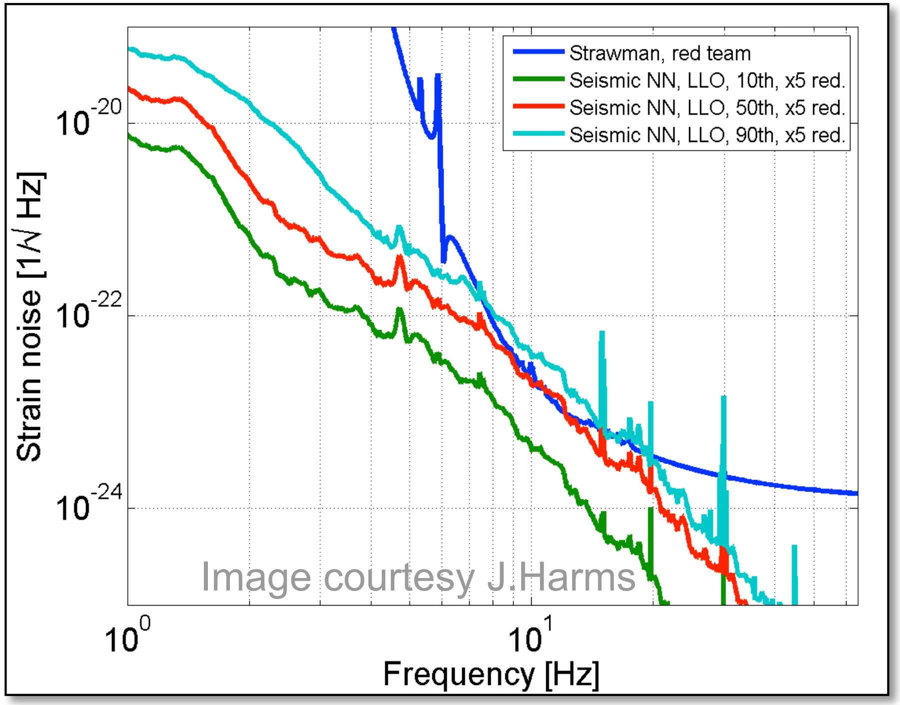
Suspension Thermal Noise



Gravity Gradient Noise

- **Red design assumes a reduction factor of 5.**
- Please note seismic noise is not constant. The factor 5 assumed guarantees that 90% of the time the Newtonian noise would be below the LIGO-3 red sensitivity.

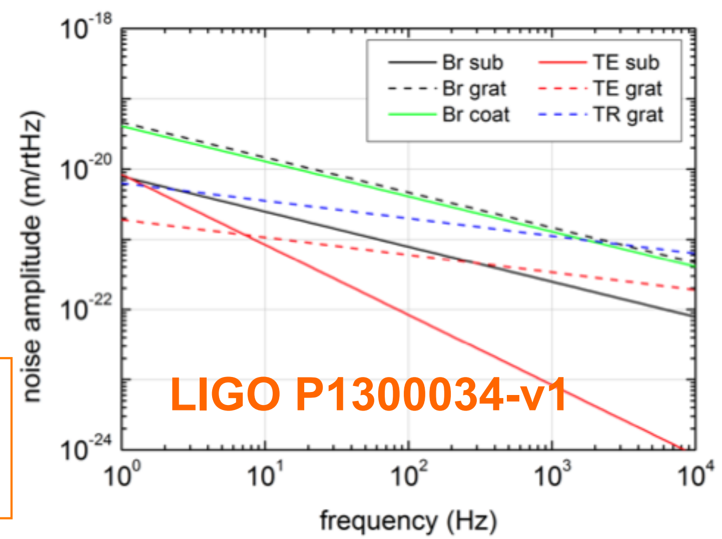




Coating Brownian noise

- Assumed an overall improvement by a factor **3.2**.
- Factor **1.6** from increased beam sizes.
- Another factor of **2** on top of this from either:
  - Better coatings
  - Khalili cavities
  - ~~Resonant waveguide mirrors~~

Recently learned that waveguides at 1064nm and 290K have higher thermal noise than a standard coating. ☹



Calculation of thermal noise in grating reflectors

D. Heinert,[1] S. Kroker,[2] D. Friedrich,[3] S. Hild,[4] E.-B. Kley,[2,5] S. Leavey,[4] I. W. Martin,[4] R. Nawrodt,[1] A. Tünnermann,[2,5] S. P. Vyatchanin,[6] and K. Yamamoto[3]

Quantum noise

- We kept the interferometer configuration and the mirror reflectivities the same as in aLIGO baseline.
- **Introduced frequency dependent input squeezing.**
- Key aspects: achievable squeezing level & required length of filter cavity

Laser and Optical Parameters	*aLIGO baseline*	*LIGO-3 red*
Laser Wavelength	1064 nm	1064 nm
Optical Power at Test Masses	730 kW	730 kW
Arm Cavity Finesse	450	450
Signal Recycling	$T = 20\,\%$, tuned	$T = 20\,\%$, tuned
Squeezing Factor	n.a.	20 dB
Filtercavity (FC) length	n.a.	300 m
FC Detuning	n.a.	-16.8 Hz
FC Input Mirror Transmittance	n.a.	425 ppm
Squeezing Losses	n.a.	$9\,\% + 30$ ppm roundtrip in FC

Squeezing losses

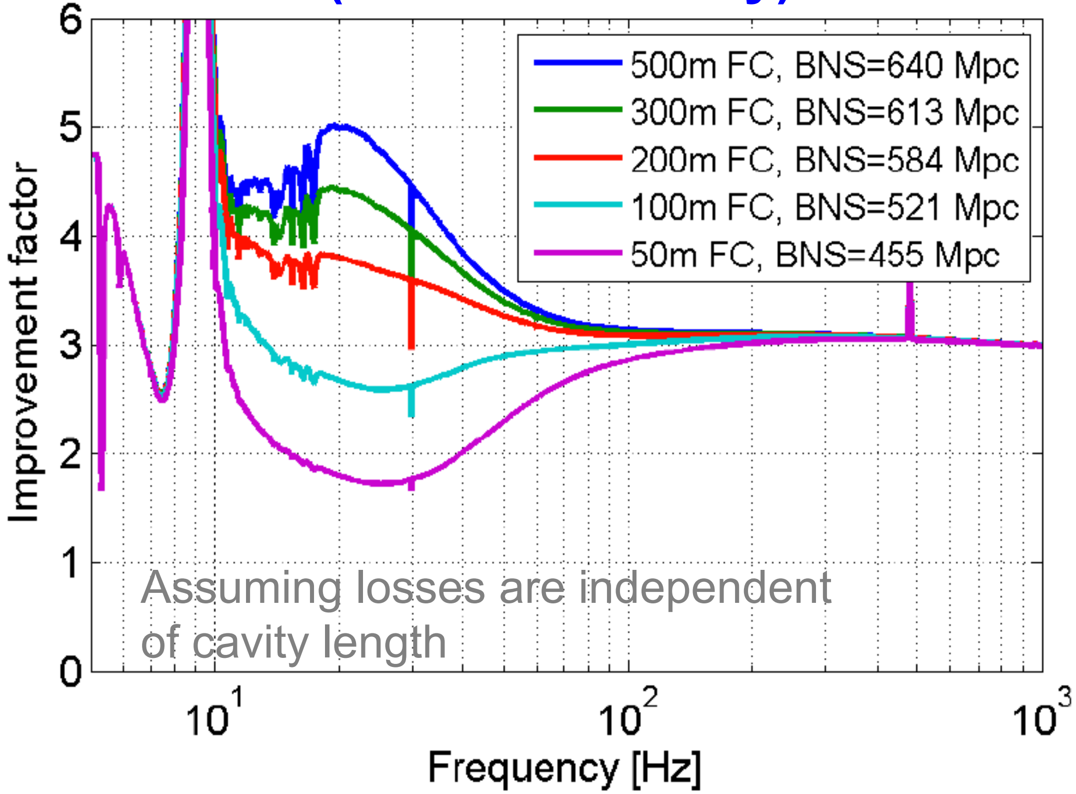
Frequency independent losses:

- Generation of squeezing: 3 %
- Optical isolation: 3 x 0.8 %
- Mode matching to IFO and to OMC: 2 x 1 %
- OMC loss and QE of PD: 2 x 0.5 %
- Mode matching to filter cavity: 1 %

= 9% in total

+

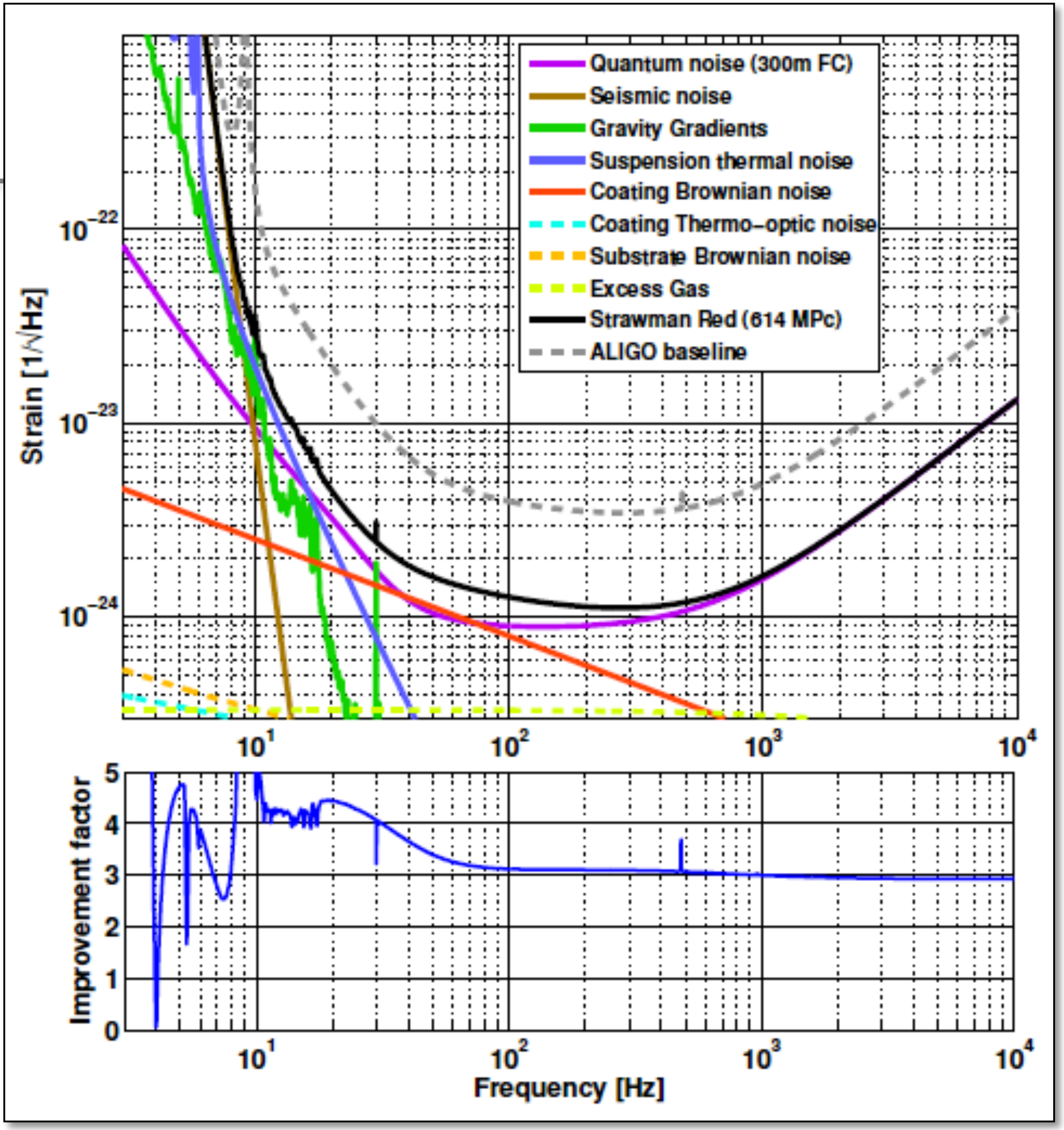
Frequency dependent loss (from filtercavity):



Starting from 20dB squeezing inside the squeezing crystal the losses reduce the observed squeezing to about 9-10dB

Team Red Sensitivity

- So if we put all the afore mentioned things together we get the following sensitivity:
- Overall an improvement of a factor 3 at all frequencies above 100 Hz. And a factor 3-4 below 100Hz.
- The binary neutron star inspiral range would improve from about 200 Mpc to above 600 Mpc.



Team Red parameters

- Rough cost estimate (only hardware included) is about 20 million $ per interferometer.
- Description of the Team Red Design can be found at https://dcc.ligo.org/cgi-bin/private/DocDB/ShowDocument?docid=78100 or docid=86550
- The sensitivity data for the Team Red design are available at https://dcc.ligo.org/cgi-bin/private/DocDB/ShowDocument?docid=86562

Strawman Red Design Overview		
Subsystem and Parameters	Advanced LIGO Baseline Design	Strawman Red Design
Sensitivity		
Binary Neutron Star Inspiral Range	200 Mpc	614 Mpc
Anticipated Strain Sensitivity	$3.5 \cdot 10^{-24}/\sqrt{\text{Hz}}$ @ 300 Hz	$1.2 \cdot 10^{-24}/\sqrt{\text{Hz}}$ @ 250 Hz
Instrument Topology		
Interferometer	Dual-recycled Michelson with Armcavities	Dual-recycled Michelson with Armcavities
Quantum Noise Reduction	n.a	Frequency-dependent input squeezing
Laser and Optical Parameters		
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Coating Noise Reduction		
Improvement Factors	n.a.	factor 1.6 from increased beam size **PLUS** factor 2 from either (i) better coatings, **OR** (ii) Khalili cavities, **OR** (iii) waveguides
Operation Temperature	290 K	290 K
IM/EM ROC	1934/2245 m	1849/2173 m
IM/EM spotsize	5.31/6.21 cm	8.46/9.95 cm
Khalili cavity length	n.a.	50 m
Gravity Gradient Noise		
Assumed Seismic Level	???	LLO ETMX, 90th percentile
Assumed subtraction factor	n.a.	5

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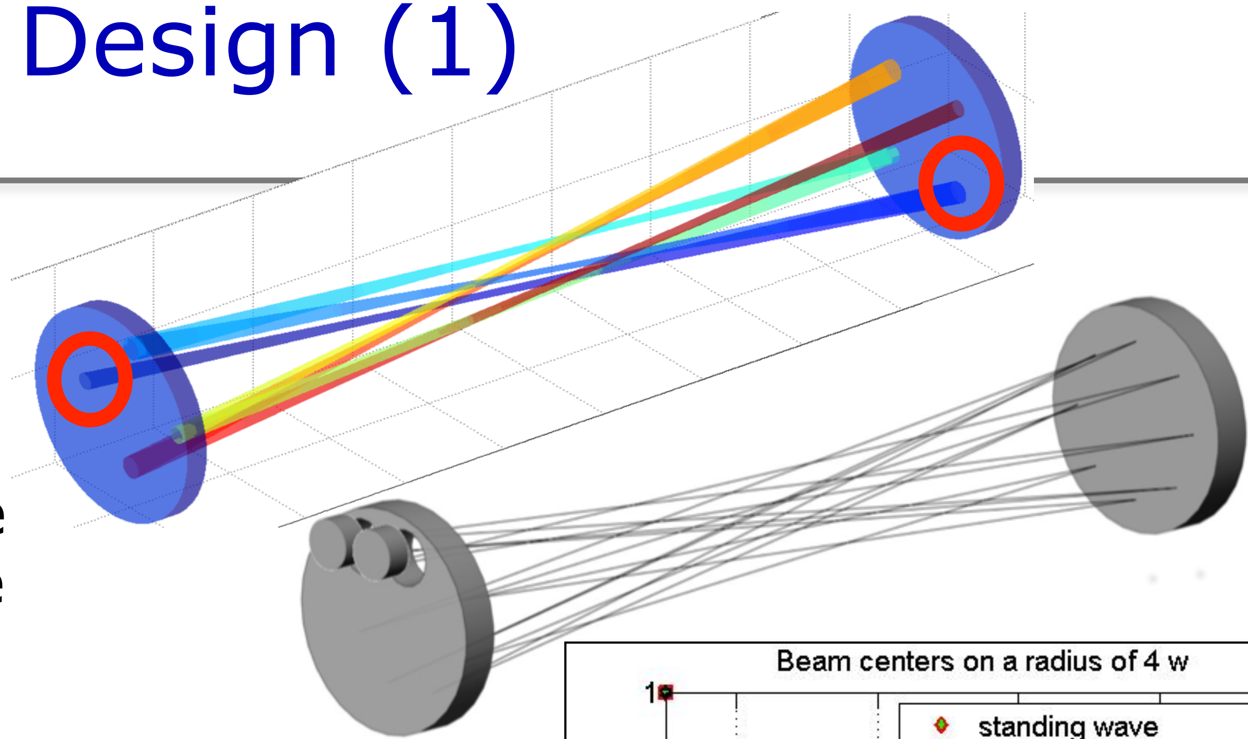
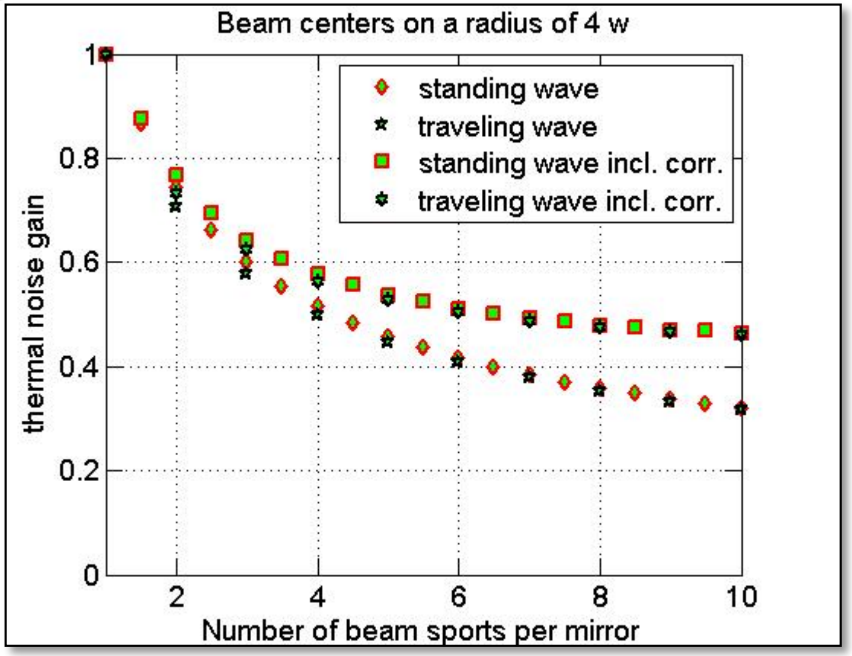
Increasing Effort, Cost, Ambition

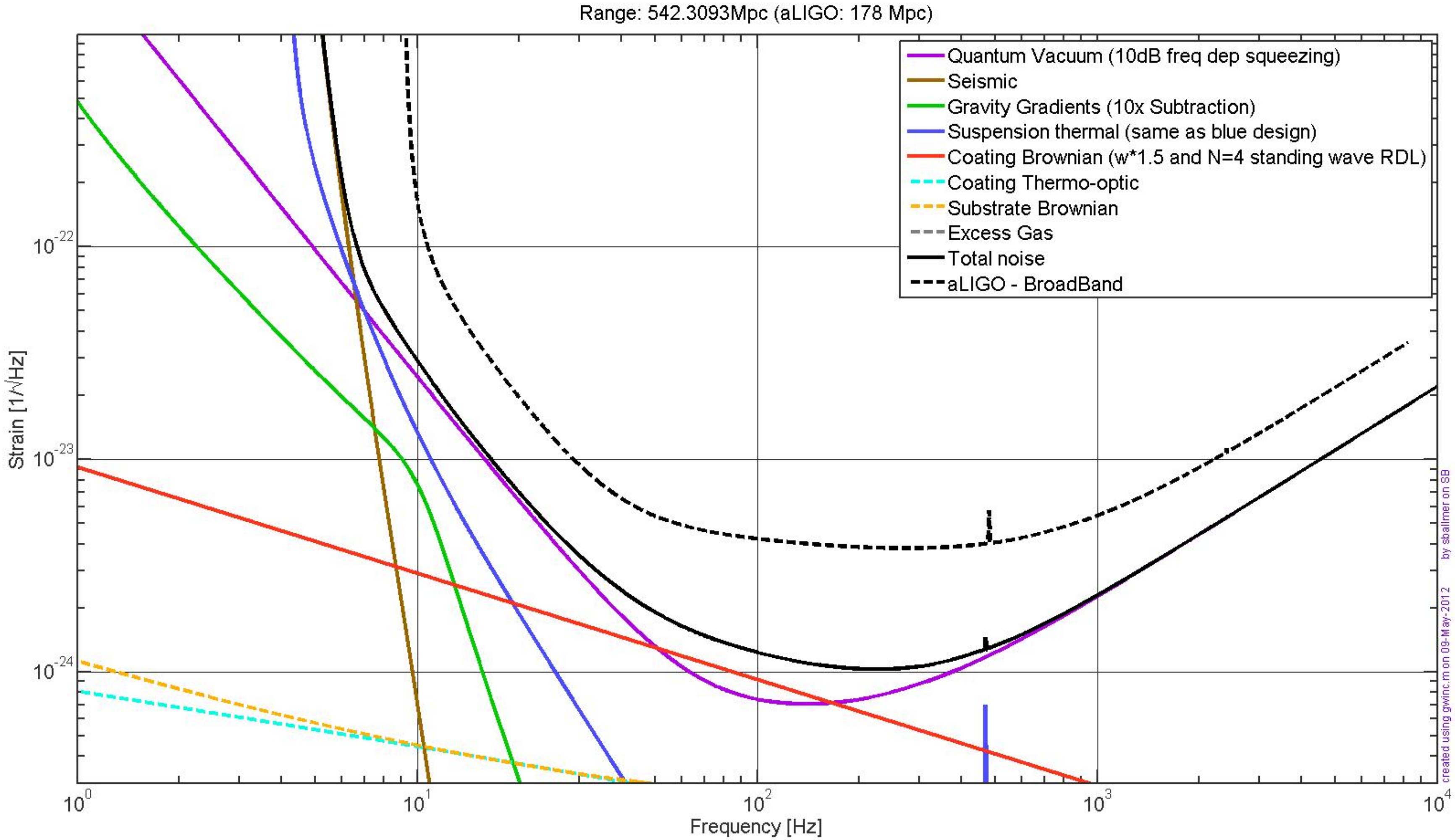
Green Design (1)

- Green designs focus on coating noise reduction.
- Idea to use a mixture of a Heriott delay line and a FB-cavity.
- 2 different modes: Traveling Wave vs Standing Wave
- Requires very large mirrors, with holes or locally modified surface



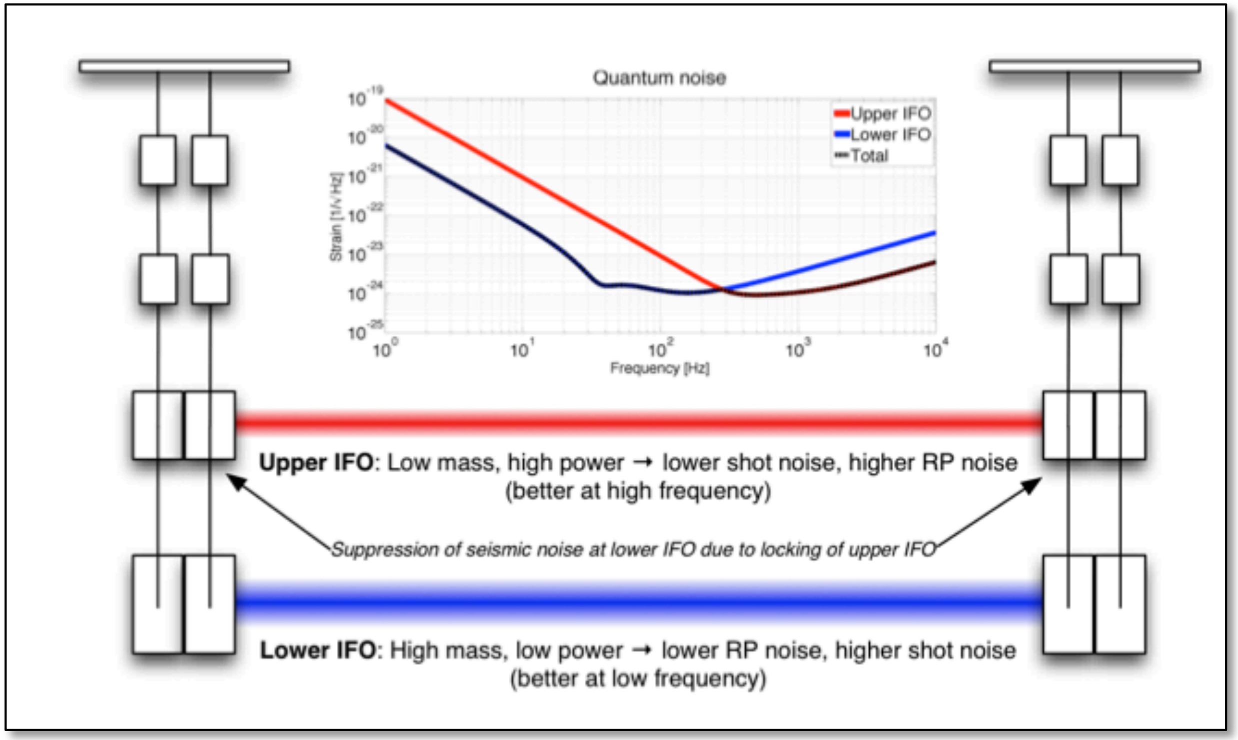


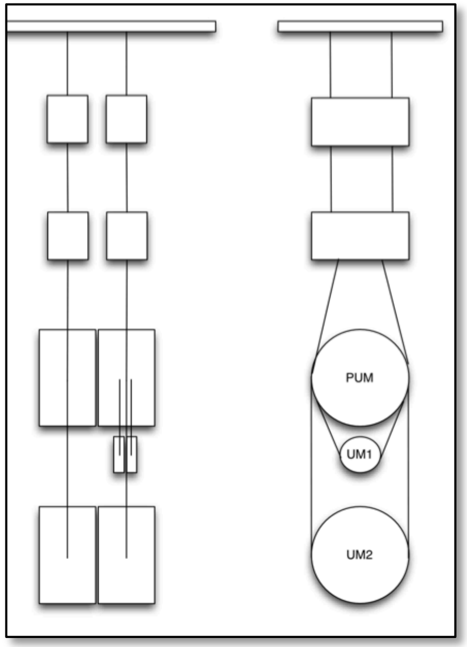
Green Design (2)



Green Design (3)

- 2nd idea investigated: Combination of a Suspension Point Interferometer + full room temperature Xylophone.





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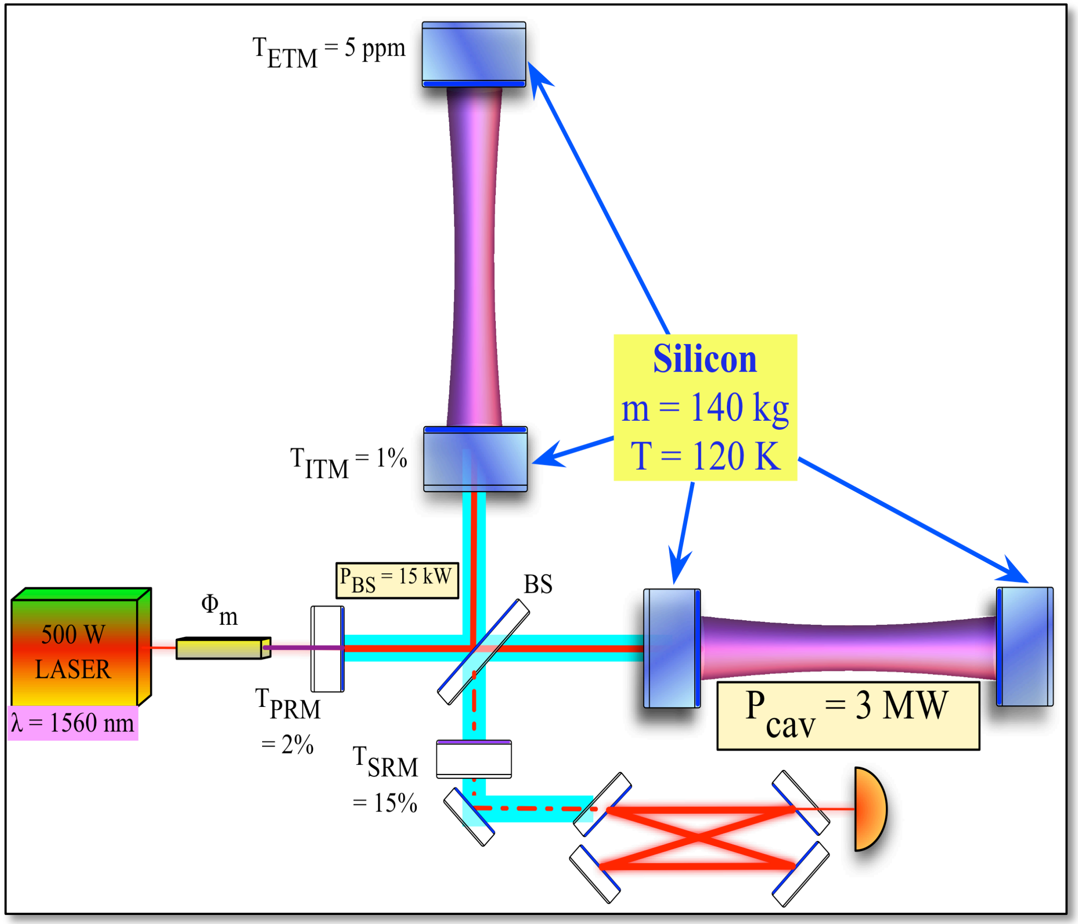
Increasing Effort, Cost, Ambition

Blue Design (1)

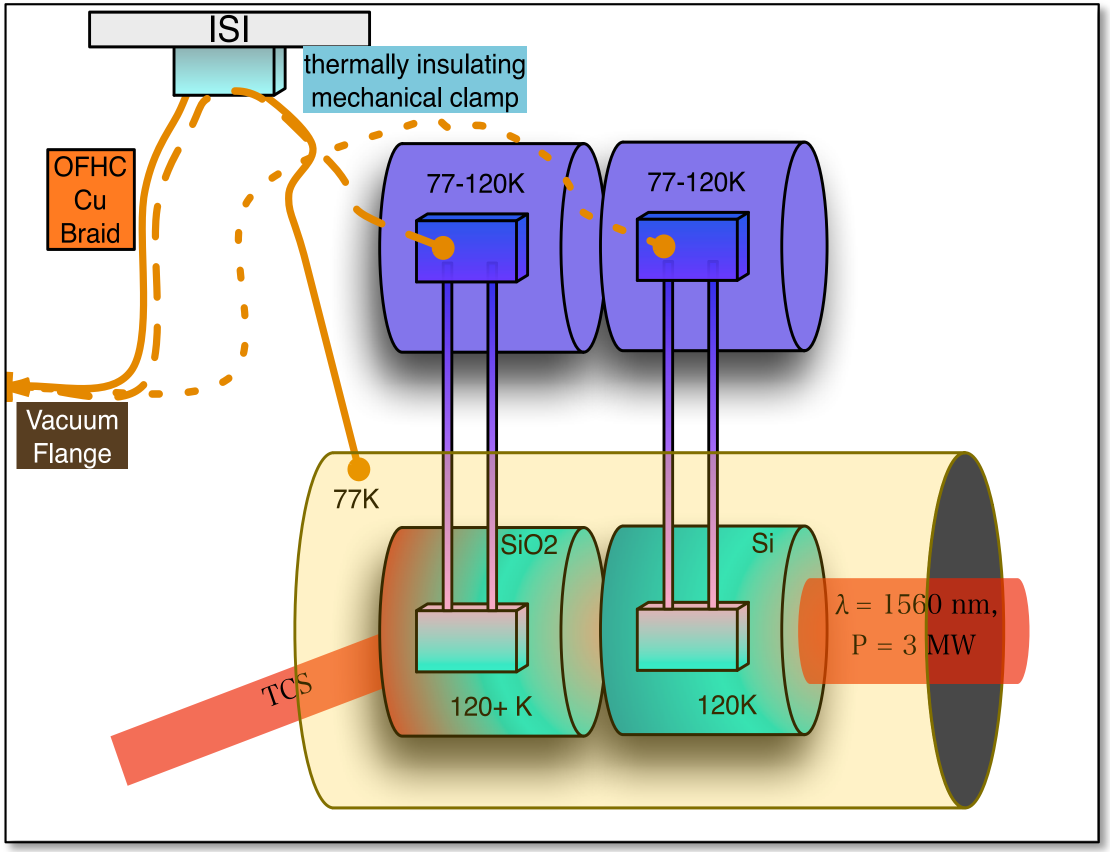
- Blue design is much more radical than red design.
- Based on cryogenic (120K) silicon test masses and suspensions to reduce thermal noise.
- Good properties of silicon:
  - Thermal expansion has a zero-crossing at 120K.
  - High thermal conductivity => smaller thermal gradients.

  Winkler et al, Phys. Rev. A 44, 7022–7036 (1991)
- Plan to use 4 times higher optical power than aLIGO.

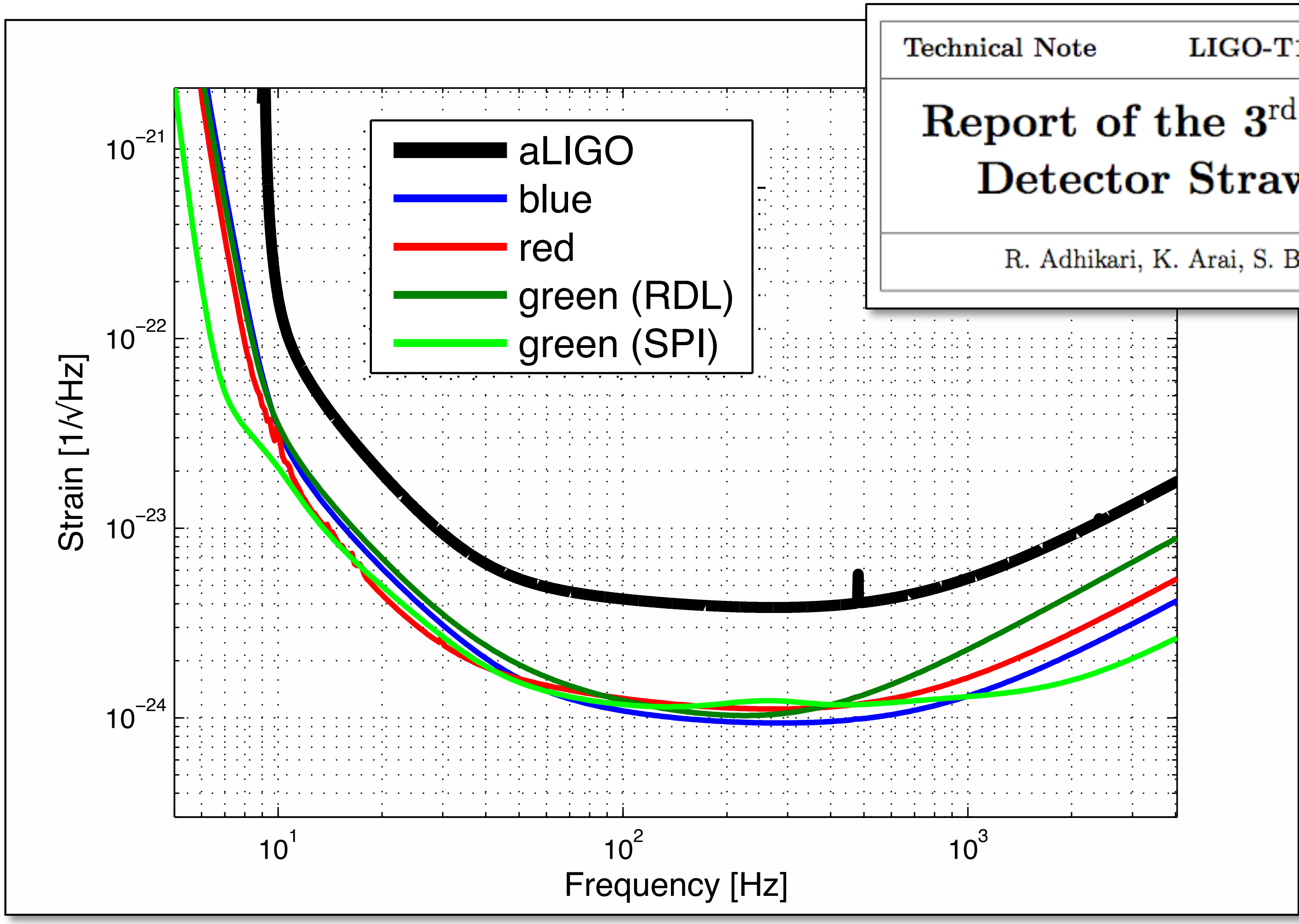


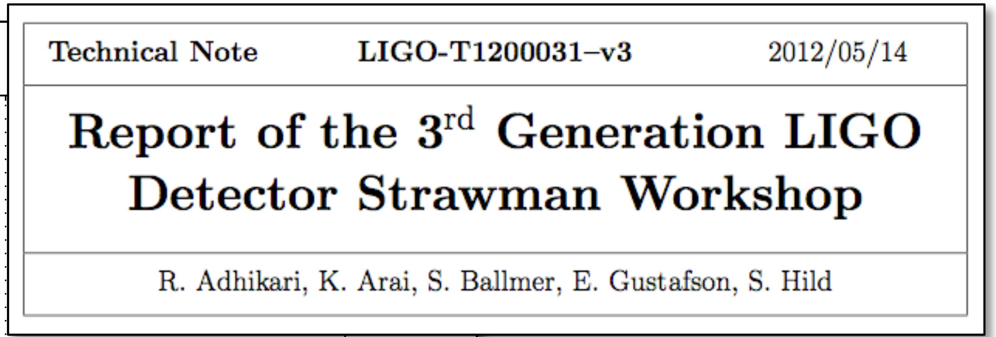
Blue Design (2)

- In contrast to the Einstein Telescope and KAGRA (both operating at 10-40K range) the cooling in the blue design will mainly be done via radiation (and not via conduction through the fibres).
- As a result the cryogenic implementation is simpler and higher optical powers can be possible.
- Lots of R&D required.
- Blue design not incremental.



How do red, blue and green designs compare in sensitivity?





Technical Note LIGO–T1200031–v3 2012/05/14

Report of the 3rd Generation LIGO Detector Strawman Workshop

R. Adhikari, K. Arai, S. Ballmer, E. Gustafson, S. Hild

Details on all design options can be found in LIGO-T1200031.

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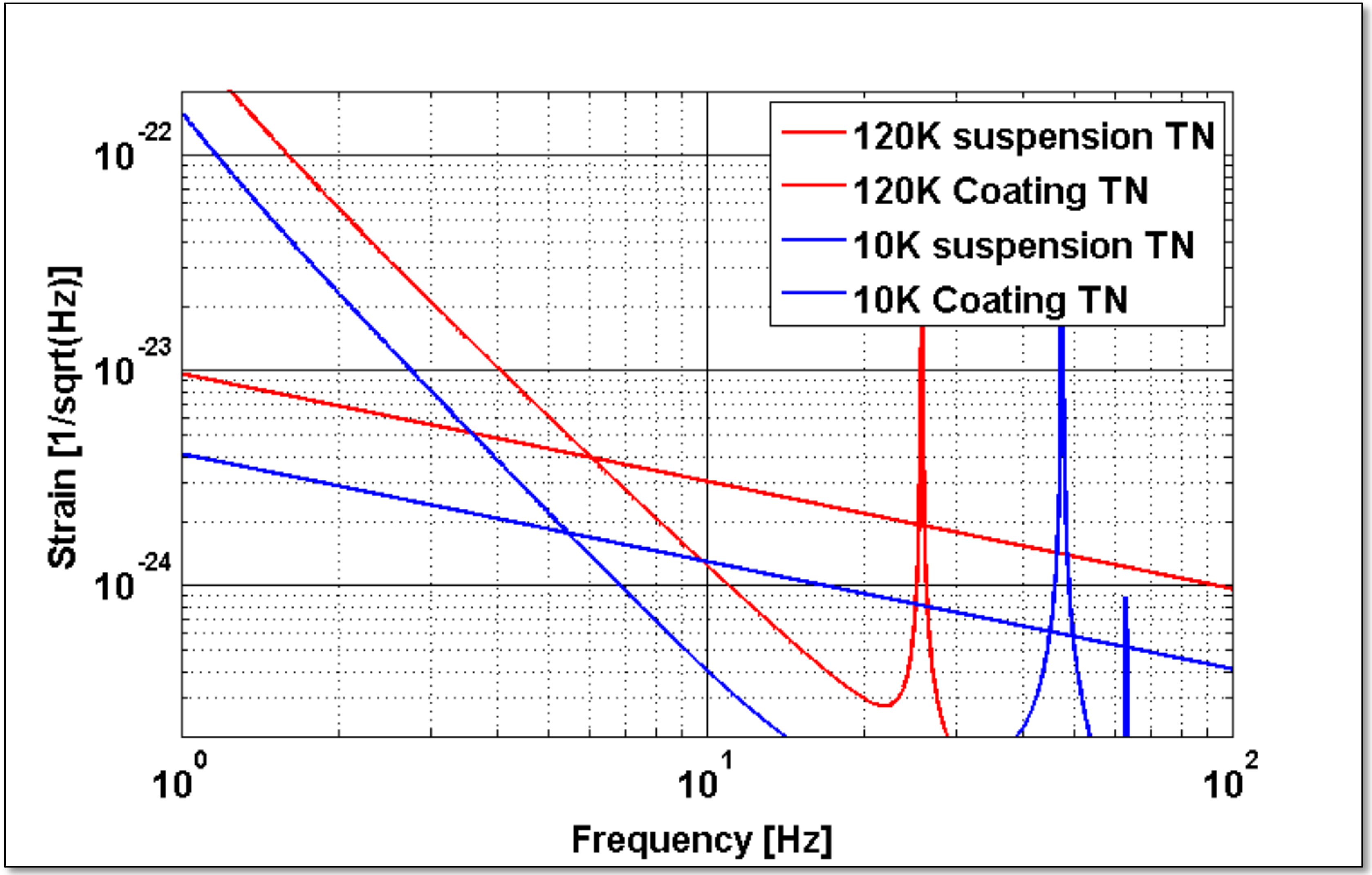
The cooler the better the noise!



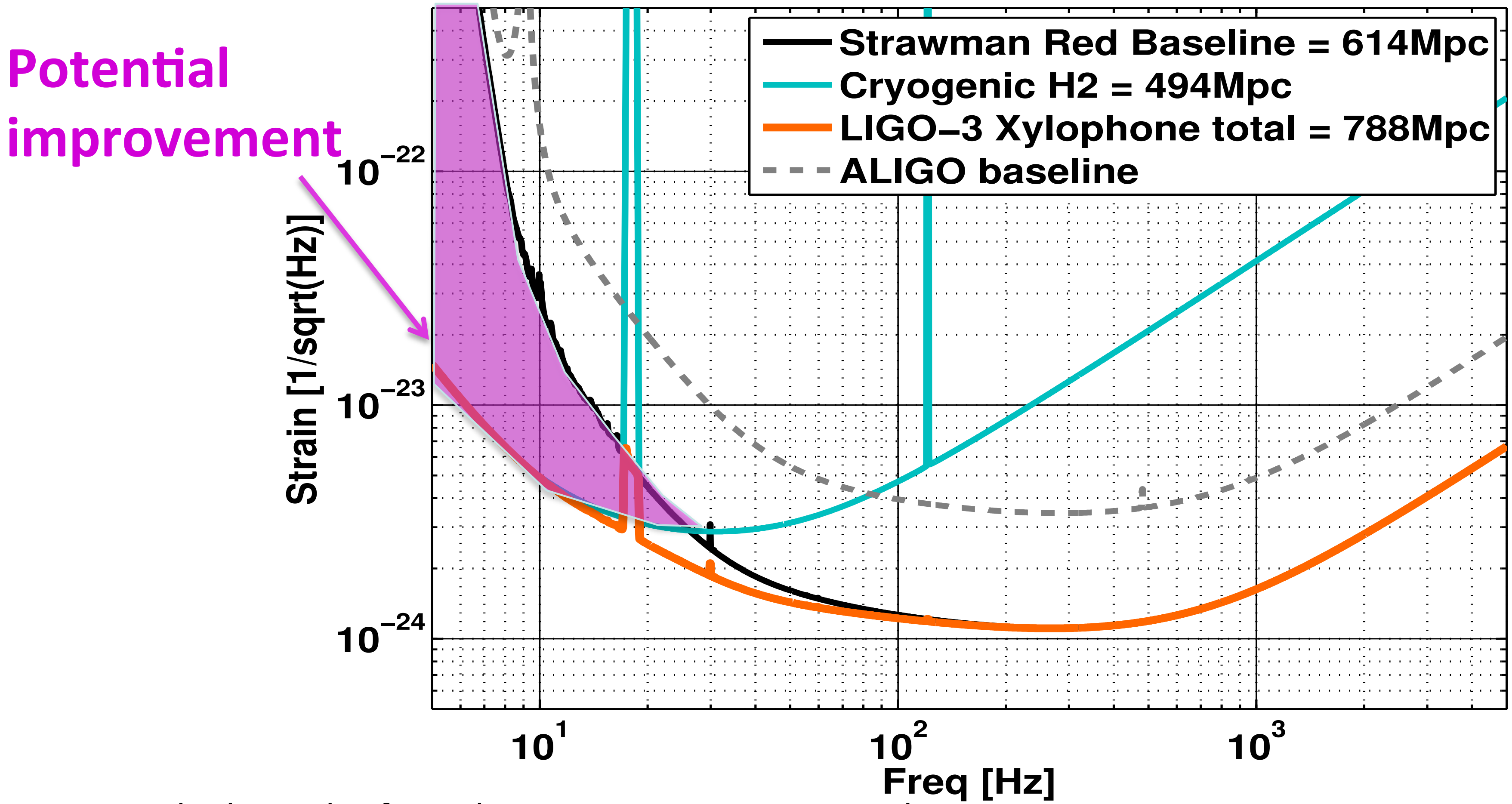
Assumes 160kg silicon test masses. Minimal fibre diameter / ribbon thickness is given by the required heat extraction for the 10K scenario and by the tensile strength for the 120K case.

Xylophone concept

- Xylophone approach: Cover the full desired frequency range by building two different interferometers, one covering the low frequency range and one covering the high frequency range.
- Resolves the problem of noise sources scaling in opposite direction (e.g. shot noise versus radiation pressure noise).
- Resolves problem of high power laser beams on cryogenic test masses.
- Please note: It is already quite amazing that our detectors can span a detection band of 2 to 3 decades in frequency.
- However, it seems likely that at some point we will find it easier (in terms of complexity) and cheaper (in terms of cost and time) to build two simpler interferometers (each optimised for the noise sources relevant in its frequency range) rather than one extremely complex instrument (optimised for 'everything').

The full xylophone

Please note: No GGN or seismic noise or any control noises are included in the LF detector noise budget !!



Numbers given in the legend refer to binary neutron star inspiral range.
A lower cut-off frequency of 5Hz was chosen.

Xylophone discussion

- If gravity gradient noise and seismic noise can be mitigated, a **cryogenic** instrument accompanying a RT partner could make a **significant low frequency sensitivity improvement**
- Using a xylophone can allow **simplifying the accompanying room temperature upgrade** (for instance shorter suspensions, lower weight of test masses, shorter filter cavity etc)
- Going for a full xylophone can give **all the benefits of a cryogenic, low-power interferometer** to cover the low frequency range while AT THE SAME TIME give the **full benefit of a not too complex and cost efficient high-power interferometer covering the high frequency** end.
- Also gives us the possibility to learn cryogenics and prepare ourselves for the future.

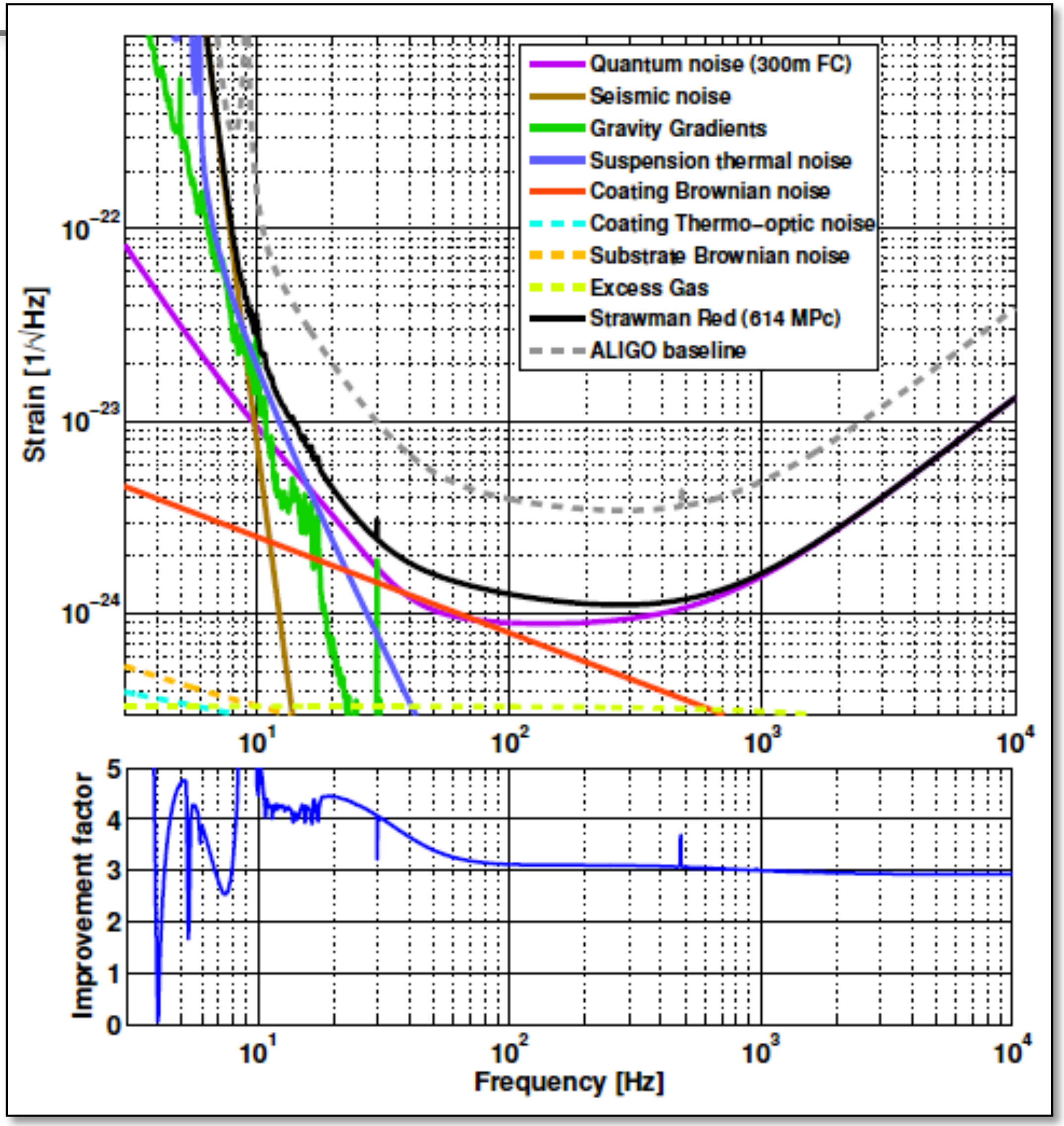
Contents

- Background / Introduction
- A Zoo of technologies to reduce individual noise sources.
- Example upgrade 1: Team Red design
- Example upgrade 2: Team Green design(s)
- Example upgrade 3: Team Blue design
- Example upgrade 4: Team Red Xylophone
- What can we learn from all this?

Increasing Effort, Cost, Ambition

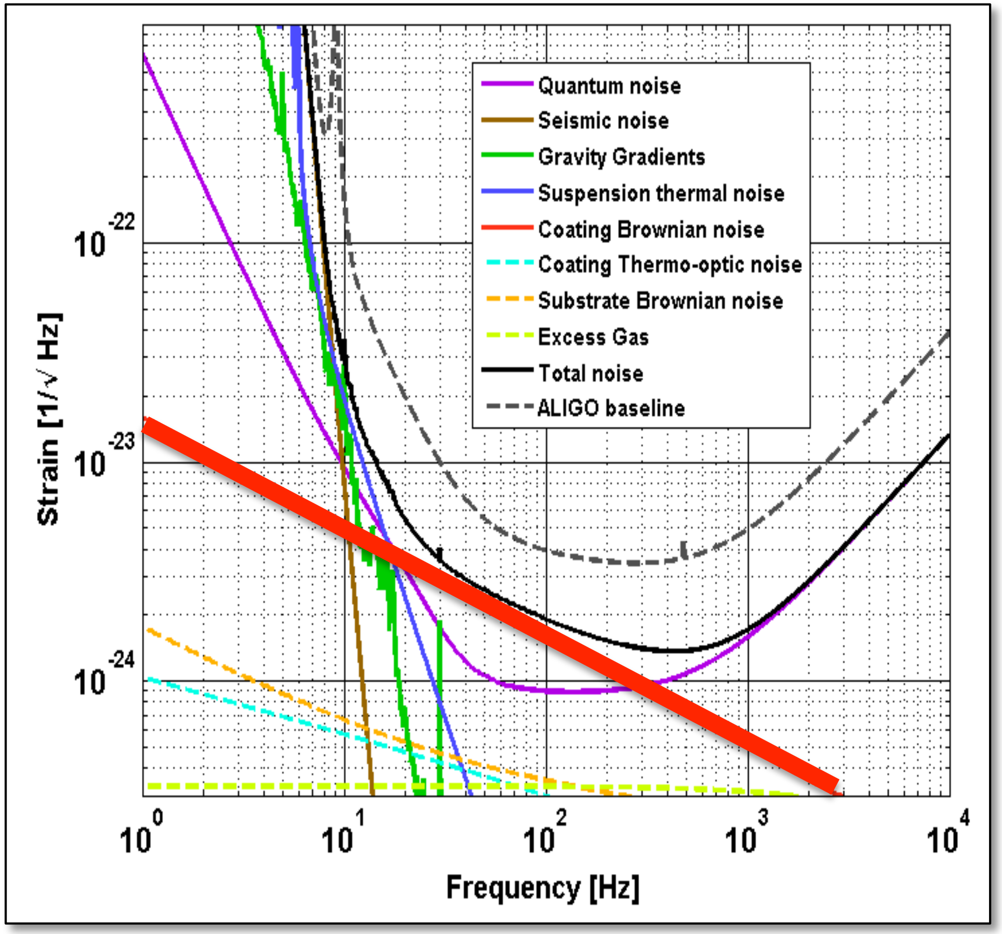
What can we learn from this? (I)

- **Advanced LIGO is far away from its facility limits.**
- The Team Red design would alllow an incremental upgrade, improving the sensitivity broadband by a factor 3-4.
- 'You can buy sensitivity at a rate of the order of about 10Mpc/$1million'.



What can we learn from this? (II)

- **If we are prepared to do without the magic factor of 2 in coating noise improvement, then**:
- We still get a substantial sensitivity improvement to a BNS range of 430Mpc.
- But even more importantly: Such a design would only include techniques and know-how that we already have! In principle we could start building such an interferometer right away.



What can we learn from this? (III)

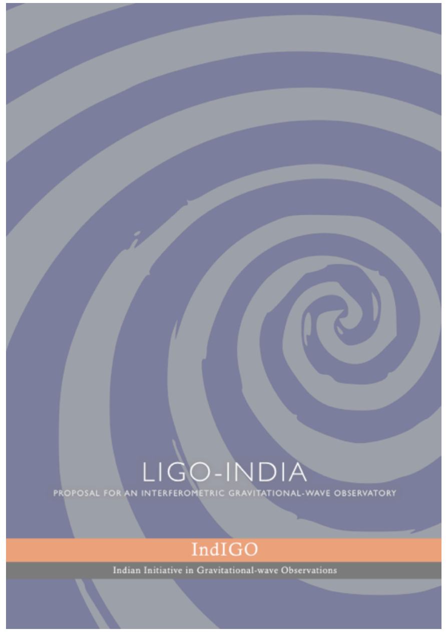
- The developed designs vary strongly and cover a wide spectrum in terms of **cost**, **technical readiness** of the **involved technologies**, the **required shut-down times** etc.
- Designs have been extremely useful for **defining what R&D is required** to be carried out over the upcoming years.
- In a few years, when required timelines and available budget are clearer as well as open R&D questions have been answered, the upgrade plans will be narrowed down to a single design.
- **Key message for the moment: There will be significant sensitivity improvements possible after Advanced LIGO will have accomplished its mission!**

R&D required for the different designs

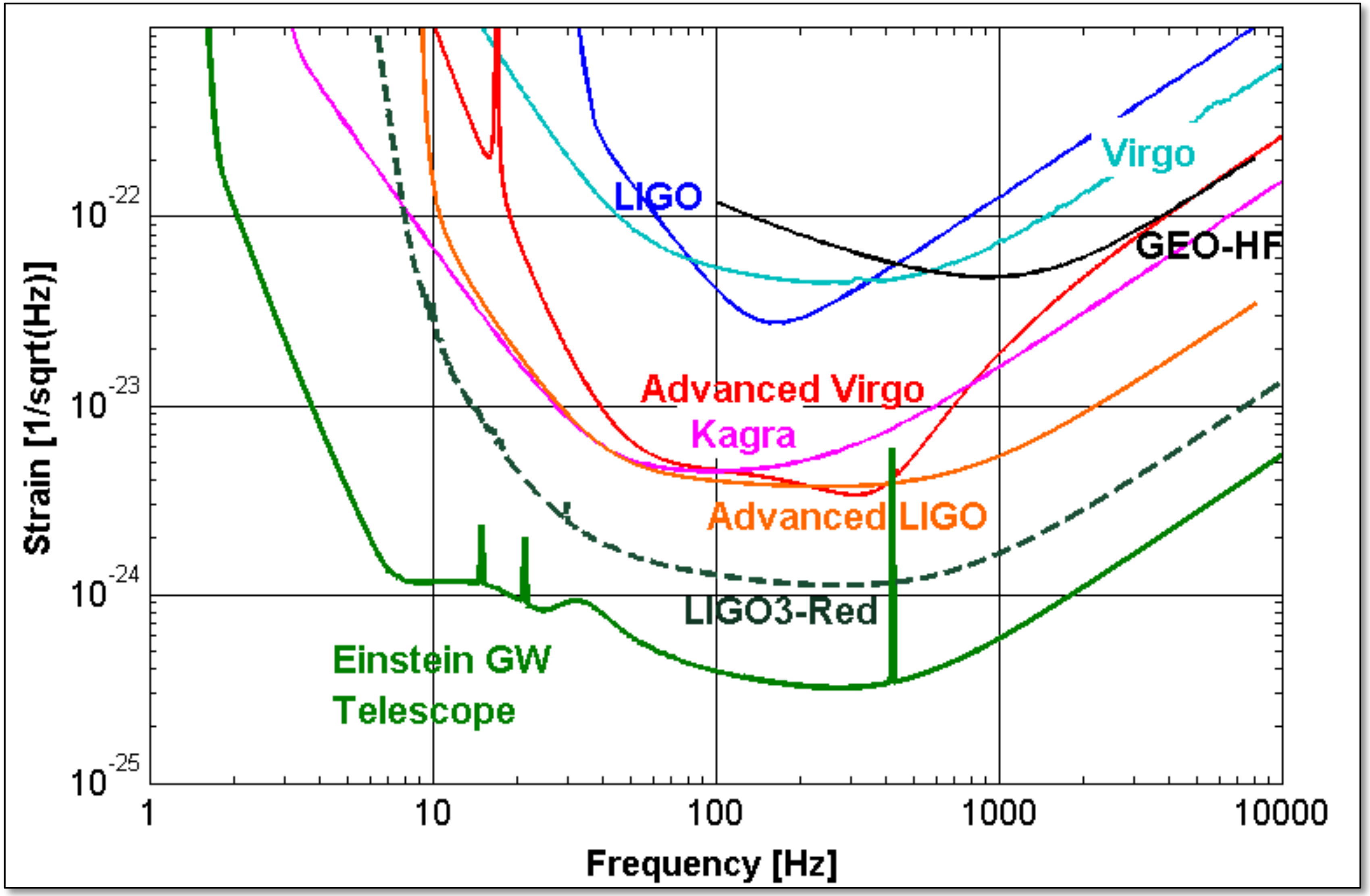
Red • Metrology for 55cm mirrors • Better coatings at 1064nm and RT • Suspensions (160kg, 1.2m)		**Green** • Metrology for ~1m mirrors • Mirrors with insets
	• Freq-dependend Squeezing • Newtonian noise subtraction	
• Silicon properties at 120K • 500 W laser at 1550nm • Infrastructure for cooling to 120K (scattering) **Blue**	• Silicon suspensions	• Silicon properties at 10-40K • Heat extraction • Thermal shields **Red Xylo**

What can we learn from this for LIGO-India?

- Now is a good time to start looking at what of these upgrades can be implemented into LIGO India right from the beginning.
- Potentially can save money if certain infrastructures are included already now.
- **Need to be ultra low-risk!**
- Potential example: Include frequency dependent squeezing.
  - Need to include vacuum tube for filter cavity
  - Satisfies ultra-low risk requirement => can easily just block the injection path to get back the baseline configuration.
- An other obvious examples? Newtonian noise reduction? ….



LIGO-3 sensitivity in context



Thanks very much for your attention...

... and we are looking forward to hear feedback what is wrong with these designs or why you think they won’t work!